



Early Journal Content on JSTOR, Free to Anyone in the World

This article is one of nearly 500,000 scholarly works digitized and made freely available to everyone in the world by JSTOR.

Known as the Early Journal Content, this set of works include research articles, news, letters, and other writings published in more than 200 of the oldest leading academic journals. The works date from the mid-seventeenth to the early twentieth centuries.

We encourage people to read and share the Early Journal Content openly and to tell others that this resource exists. People may post this content online or redistribute in any way for non-commercial purposes.

Read more about Early Journal Content at <http://about.jstor.org/participate-jstor/individuals/early-journal-content>.

JSTOR is a digital library of academic journals, books, and primary source objects. JSTOR helps people discover, use, and build upon a wide range of content through a powerful research and teaching platform, and preserves this content for future generations. JSTOR is part of ITHAKA, a not-for-profit organization that also includes Ithaka S+R and Portico. For more information about JSTOR, please contact support@jstor.org.

NOTE ON THE BINARY STARS β 208 AND β 524.

Two recent observations of the binary star β 208 made with the 36-inch telescope show that the angular motion of the companion-star has been fully 90° since 1898. The distance has diminished decidedly in the same interval. My measures give the following position:—

$$1905.18 \quad 167^\circ.6 \quad 0''.30 \quad 2^{\text{h}}.$$

If, as now seems likely, the orbit is very eccentric, it will be possible in five or six years to compute a satisfactory orbit for this pair. The system is also of interest because of its large proper motion— $0''.48$ in the direction 328° .

The binary star 20 *Persei* = β 524 has for several years been a difficult object to measure, even with the 36-inch telescope. The apparent distance between the components is now increasing, and will probably continue to increase for a number of years, though its maximum value will not greatly exceed $0''.3$. As this pair belongs to the class of short-period binaries, the periodic time not being much more than thirty years, it deserves annual measures by observers having telescopes adequate to such work.

My last measures are:—

$$1904.83 \quad 5^\circ.8 \quad 0''.16 \quad 2^{\text{h}}.$$

March 24, 1905.

R. G. AITKEN.

NOTE ON COMET *c* 1904.

Comet *c* 1904, discovered by M. BORRELLY on December 28, 1904, has proved to be an object worthy of more attention than was at first suspected, for it is traveling in an elliptic orbit, and hence is a member of the solar system, not a chance visitor.

This discovery was made by M. FAYET at Paris, and, independently, by the present writer. My orbit, based upon my observations of December 31, 1904, January 17, and January 27, 1905, gives a period of 7.3 years; M. FAYET's revised elements, derived from normal places representing the observations from December 30, 1904, to January 26, 1905, make the period a little shorter, 7.0 years.

In other respects the two orbits are very similar, and the